

# Discovery of eleven new ZZ Ceti stars<sup>\*</sup>

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**Abstract.** We report the discovery of eleven new ZZ Ceti stars using telescopes at OPD (Observatório do Pico dos Dias/LNA) in Brazil, the 4.1 m SOAR (Southern Astrophysical Research) telescope at Cerro Pachon, Chile, and the 2.1 m Otto Struve telescope at McDonald observatory. The candidates were selected from the SDSS (Sloan Digital Sky Survey) and SPY (ESO SN Ia progenitor survey), based on their  $T_{\text{eff}}$  obtained from optical spectra fitting. This selection criterion yields the highest success rate of detecting new ZZ Ceti stars, above 90% in the  $T_{\text{eff}}$  range from 12 000 to 11 000 K. We also report on a DA not observed to vary, with a  $T_{\text{eff}}$

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<sup>\*</sup> Partially based on observations at Observatório do Pico dos Dias/LNA, the Southern Astrophysical Research telescope, a collaboration between CNPq-Brazil, NOAO, UNC and MSU, and McDonald Observatory of the University of Texas at Austin.

placing the star close to the blue edge of the instability strip. Among our new pulsators, one is a little bit cooler than this star for which pulsations were not detected. Our observations are an important constraint on the location of the blue edge of the ZZ Ceti instability strip.

**Key words.** (Stars): white dwarfs, Stars: variables: general, Stars: oscillations

## 1. Introduction

White dwarf stars are the end point of the evolution of 95-98% of all stars. As they cool, white dwarf stars pass through three distinct instability strips, depending on their temperatures, atmospheric composition, and the element that drives pulsation: carbon and/or oxygen in DOVs, helium in DBVs, or hydrogen in DAVs or ZZ Ceti. Observationally, the DA instability strip ranges in effective temperature ( $T_{\text{eff}}$ ) from 12 270 to 10 850 K (Bergeron et al. 2004, Mukadam et al. 2004, and Gianninas et al. 2005). While Bergeron’s instability strip is pure, Mukadam et al. (2004) and Mullally et al. (2005) did not detect light variability in a significant number of stars located within the  $T_{\text{eff}}$  and  $\log g$  boundaries of the instability strip. If the instability strip is contaminated (occupied by non-pulsators of the same spectral type as pulsators), pulsation would no longer be a normal stage of white dwarf evolution. However, because of destructive interference between modes, some white dwarf stars can appear to be constant for timescales of 3-4 hours. Another important issue is that some known pulsating white dwarf stars show amplitudes as low as 4 mma (e.g. Kanaan et al. 2005), so we have to reach *at least* this limit to determine whether a star pulsates or not; many stars not observed to vary by Mukadam et al. 2004 have detection limits above 4 mma.

Pulsations are global; each periodicity is an independent measurement of the interior. This gives rise to asteroseismology: the study of stars below the photosphere through measurement of their pulsations. Comparing the observed pulsation spectrum to models of the stellar interior through asteroseismology allows for the stellar mass determination (Winget et al. 1990, Bradley & Winget 1994) and even constrain the value of the  $\text{C}^{12}(\alpha, \gamma)\text{O}^{16}$  cross section (e.g. Metcalfe et al. 2002), which can only be measured in a terrestrial laboratory at energy levels eight orders of magnitude higher. Furthermore, the cooling time scales of DAVs (Kepler et al. 2000, 2005, Mukadam et al. 2003) can be used to calibrate the age of the galactic morphological components (Winget et al. 1987, Hansen et al. 2002, von Hippel 2005), by observing field stars, open and globular clusters up to magnitudes to include the turnoff of the white dwarf cooling sequence.

Pulsating white dwarf stars are also important for the study of extreme physics: internal crystallization (Winget et al. 1997, Kanaan et al. 2005), neutrino cooling (Kawaler et al. 1986, Winget et al. 2004) and axion emission (Córscico et al. 2001, Kepler 2004,

Kim et al. 2005). Both particles can be created at temperatures and densities found in white dwarf cores and consequently must be taken into account in the cooling models.

Quasar surveys, like SDSS (Sloan Digital Sky Survey), HE (Hamburg ESO) and 2dF (Two Degree Field), are increasing the number of spectroscopically identified white dwarf stars as a by-product; quasars and white dwarf stars have similar colors. As these new white dwarf stars are fainter than the ones previously known, but with similar  $T_{\text{eff}}$ , we are able to study a sample of stars at larger distances, with a wider range of progenitor metallicity. The study of their chemical composition is of interest if we are to use SN Ia as standard candles, that are products of mass accretion onto white dwarf stars.

Mukadam et al. (2004) show that at least 90% of the SDSS candidates with  $12000 \geq T_{\text{eff}} \geq 11000$  K are pulsators. We selected our targets based on the  $T_{\text{eff}}$  we derive by fitting the whole optical spectra, as described by Kleinman et al. (2004). For the candidates with spectra obtained by the SPY survey (e.g. Napiwotzki et al. 2003), we derive  $T_{\text{eff}}$  by fitting only the hydrogen line profiles. In both fitting procedures, we use the same model atmosphere grid, developed by Detlev Koester (similar to that described in Finley et al. 1997). We only obtained time series photometry of stars with  $T_{\text{eff}}$  inside or close to the ZZ Ceti instability strip. All observed stars, but one, turned out to be pulsators.

In this paper, we report the discovery of eleven new pulsators and one star not observed to vary down to a detection limit of 3.75 mma.

## 2. Observations

We observed some of our candidates at Observatório Pico dos Dias, LNA, in Brazil. For these observations, we used the 1.6-m telescope, with a frame transfer CCD 301, focal reducer, and no filters. The CCD has a quantum efficiency of 60% at around  $4000\text{\AA}$ , the preferred wavelength range to observe ZZ Ceti to maximize the pulsation amplitude versus CCD efficiency. The integration times range from 15 to 45 seconds, depending on readout times, weather conditions, and apparent stellar magnitude. We observed without filters because the g-mode pulsations in ZZ Ceti stars are coherent at all optical wavelengths (Robinson, Kepler & Nather 1982). We also observed bright candidates with the 0.6-m Zeiss telescope. We used the CCD 106, which has a 7.5 s readout, and integrations of 30 to 60 seconds.

In June and July 2005, we used the SOAR Optical Imager, a mosaic of two EEV  $2048 \times 4096$  CCDs, thinned and back illuminated, with an efficiency around 73% at  $4000\text{\AA}$ , at the cassegrain focus of the 4.1-m SOAR telescope. We observed in fast readout mode, with the CCDs binned  $4 \times 4$ , to decrease the readout+write time to 6.4 s, and still achieve a  $0.354''/\text{pixel}$  resolution. The exposure times ranged from 20 to 40 s. All observations were obtained with a Johnson B filter.

We also observed with the Otto Struve 2.1-m telescope at McDonald Observatory, using the Argos camera, a frame transfer CCD in the prime focus, with a BG40 filter to reduce the scatter from sky (Nather & Mukadam 2004). The integration times were 20 and 30 s.

In Table 1, we show the journal of observations for all stars.

The technique to detect variability is differential time series photometry, comparing the targets with the other stars in the same field, to minimize the effects of sky and transparency fluctuations.

For each run, we extracted light curves using the IRAF script `hsp`, developed by Antonio Kanaan, using different aperture sizes and select the light curve with either the lowest noise level or the highest signal-to-noise rate (SNR) in the Fourier transform.

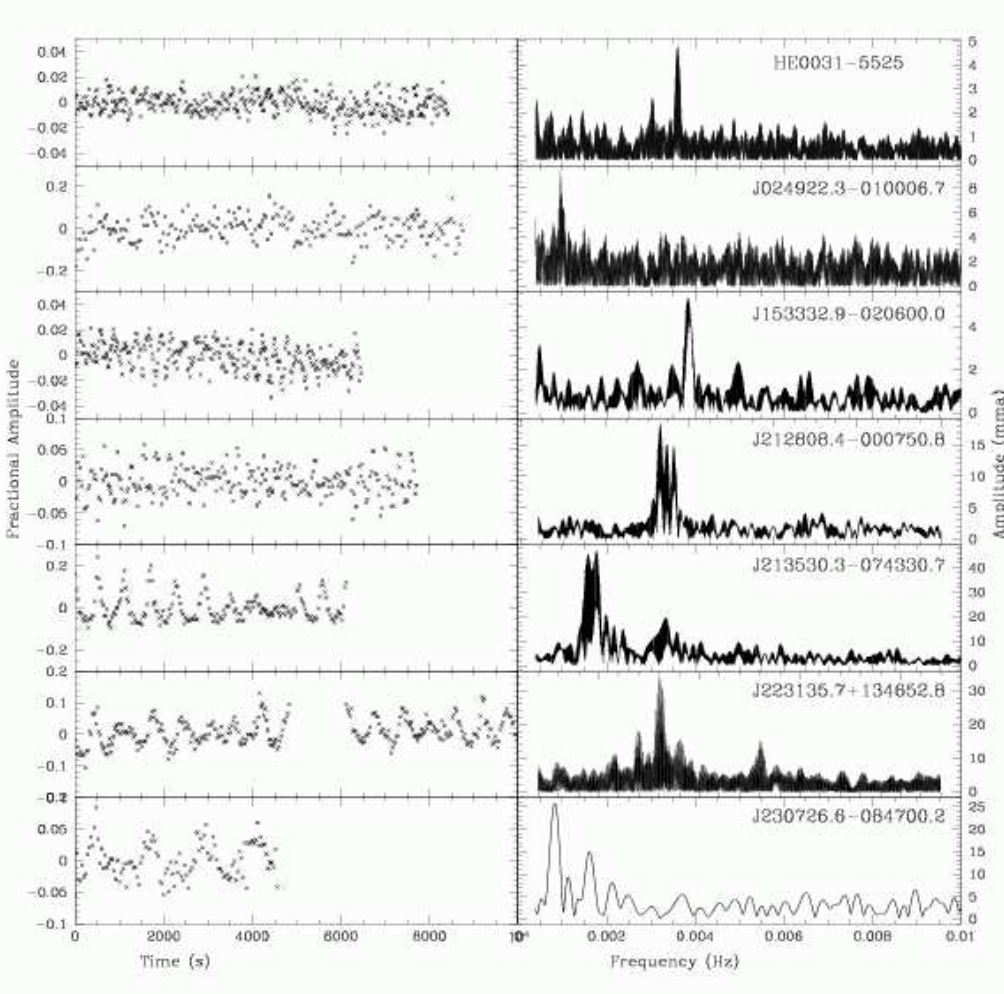
### 3. Results

In Table 2, we list the new variable stars and their physical parameters derived from optical spectra, using an improved model grid similar to that described in Finley et al. (1997). We fit the candidates from SDSS using the whole optical spectra and photometry, as in Kleinman et al. (2004), while for the candidate from HE we used only the hydrogen lines in the optical spectra, because it is independent of the flux calibration and the SNR is high enough.

In Fig. 1 and 2, we show the light curves on the left panels and the Fourier transform (FT) on the right. We show only a small part of the light curves, but the FTs were calculated for the entire dataset for each target. The lower detection limit in frequency is  $2/T$ , where  $T$  is the total length of the light curve, while the higher is the Nyquist frequency,  $f_{\text{Nyquist}} = 2/t_{\text{exp}}$ .

The criterion we used to determine which peaks are real in the discrete FT is to adopt an amplitude limit such that the probability of a noise peak exceeding this value is only 1/1000 (false alarm probability). We then repeatedly subtracted the largest signal in the FT until there were no more peaks above the detection limit (see Table 3). The uncertainties in frequency are, on average,  $123\mu\text{Hz}$ , equivalent to 1s at  $P \sim 100\text{s}$ .

Combining these eleven new ZZ Ceti with the previously known pulsators, we see a well defined instability strip in Fig 3. The triangles represent previously discovered candidates from SDSS, the circles are the stars fitted by Bergeron et al. (2004), and the squares are from this paper. It is important to point out that this is not a homogeneous sample, as the previous determinations of  $T_{\text{eff}}$  and  $\log g$  were not obtained with the same model grid (e.g. Bergeron et al. 2004, Koester & Allard 2000), but the instability strip is restricted to a narrow range of temperature  $12\,270\text{ K} \geq T_{\text{eff}} \geq 10\,850\text{ K}$ . The pulsating star J235040.72-005430.9 (Mukadam et al. 2004) is clearly outside the ZZ Ceti instability strip. Despite this star being cooler than the red edge of the instability strip,

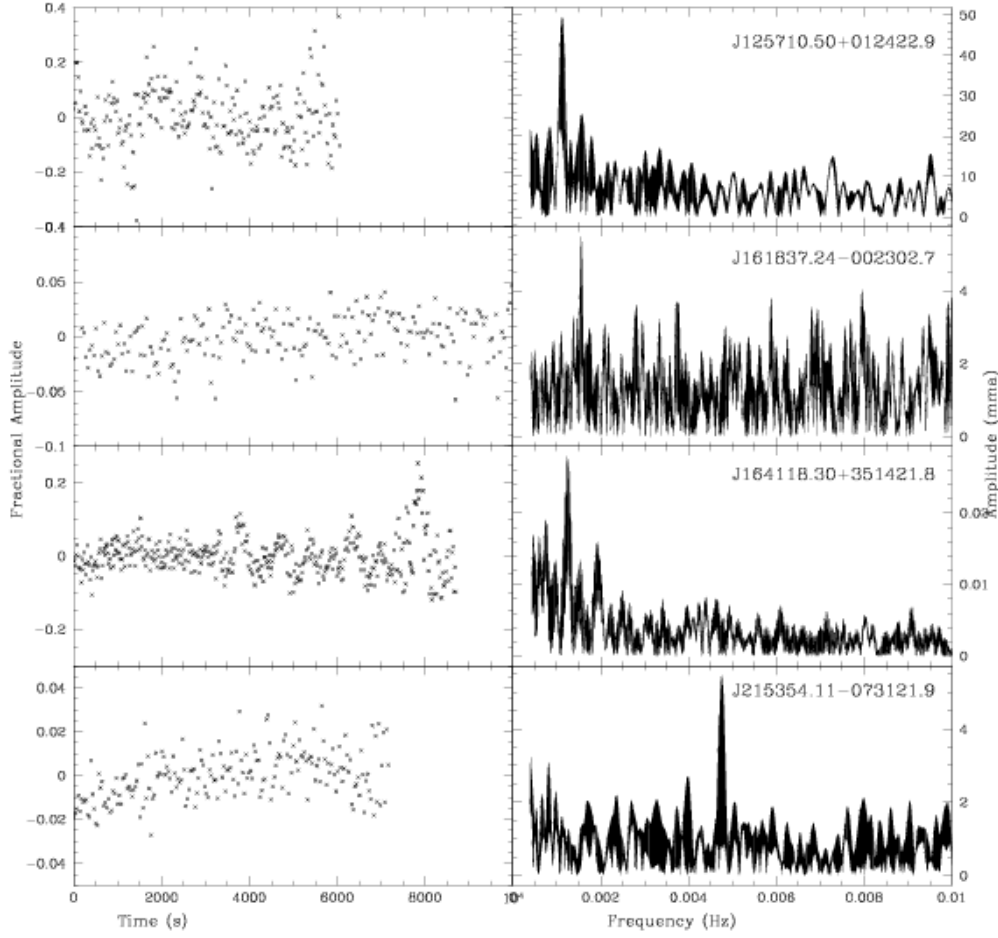


**Fig. 1.** The light curves of a single run (left panels) and the Fourier transforms of the complete data set (right panels) for the new ZZ Ceti, discovered at LNA.

its detected periods indicate it is a hot DAV. It is necessary to obtain a higher signal-to-noise spectrum, because the SDSS spectrum has  $\text{SNR} \sim 12$ , and analyze it with the line profile technique for an accurate determination of both  $T_{\text{eff}}$  and  $\log g$ . Finally, it is necessary to increase the sample of variables and non-variables at the borders of the instability strip, to determine its exact boundaries.

The star J214723.7-001358.4 did not show variability during 2 independent runs, down to a limit of 3.75 mma. According to the SDSS spectrum, this star has  $T_{\text{eff}} = 12000 \pm 280$  and  $\log g = 7.92 \pm 0.11$ , slightly hotter than the hottest ZZ Ceti from SDSS, reported in this paper, J215354.1-074330.7 (see Table 2 for details).

However, the temperatures derived from SDSS optical spectra have external uncertainties larger than 300 K, as demonstrated from their duplicate spectra; the uncertainties we quote in Table 2 are the internal uncertainties in the  $\chi^2$  fitting, assuming no correlation between  $T_{\text{eff}}$  and  $\log g$ . We conclude our observations are still consistent with a pure instability strip, but it does not exclude possible contaminations, as the strip covers only



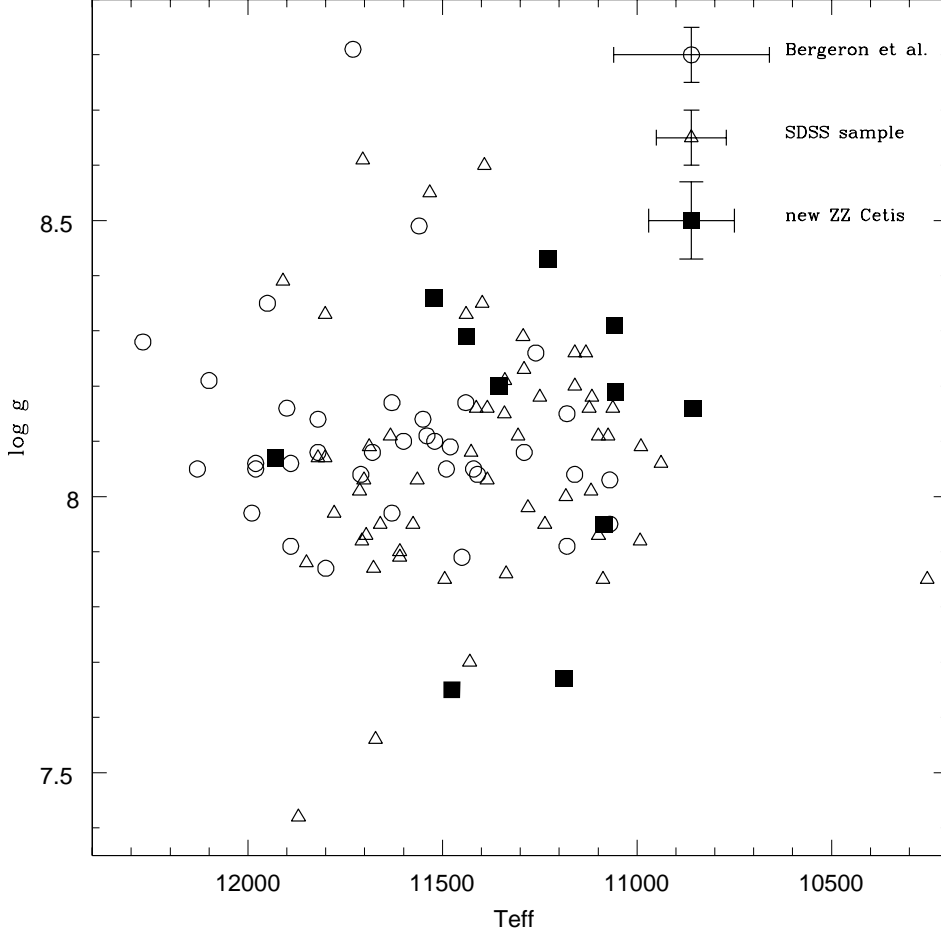
**Fig. 2.** The light curves of a single run (left panels) and the Fourier transforms of the complete data set (right panels) for the new ZZ Ceti, discovered at SOAR and McDonald observatory.

$\sim 1200$  K. To solve this problem, we need spectra with  $\text{SNR} \geq 50$  to achieve  $\sigma_{T_{\text{eff}}} \leq 200$  K and to re-observe the stars which do not show variability in the literature to decrease their detection limits below 4mma. Meanwhile, the purity of the ZZ Ceti instability strip is still an open discussion.

Our results are mapping not only the blue edge, but also the red edge of the instability strip. We discovered a pulsating star, J161837.2-002302.7, with  $T_{\text{eff}}$  and periodicities characteristics of the red edge, even though the amplitudes are significantly lower than expected. This star will be extremely important as we try to understand how the ZZ Ceti will stop pulsating around  $T_{\text{eff}} \sim 10850$  K.

#### 4. Concluding remarks

We report the discovery of eleven new pulsating DA stars, bringing the total to 106 known variables, all in the narrow temperature range,  $12\,270 \text{ K} \geq T_{\text{eff}} \geq 10\,850 \text{ K}$ , corresponding



**Fig. 3.** The updated instability strip including all known ZZ Ceti. The circles are the stars described in Bergeron et al. (2004), the triangles are the pulsators discovered by Mukadam et al. (2004), Mullally et al. (2005), and Kepler et al. (2005b) from the SDSS, and the squares are the eleven new ZZ Ceti presented in this paper. The average uncertainties for each set are on the top right of the plot.

to the partial ionization of hydrogen and development of a sub-surface convection zone. Three stars, HE 0031-5525, SDSS J153332.9-020600.0 and SDSS J212808.4-000750.8, have pulsation characteristics of hot DAVs, even though their derived  $T_{\text{eff}}$  are 11480 K, 11350 K, 11440 K, respectively. However,  $T_{\text{eff}}$  is not the only parameter which determines the position of the instability strip. Giovannini et al. (1998) demonstrate that the instability strip also depends on  $\log g$ . We therefore need to map the instability strip in both  $T_{\text{eff}}$  and  $\log g$  using accurate determinations of these parameters for the new variables.

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Star	Run start (UT)	$t_{\text{exp}}$ (s)	$\Delta T$ (hr)	# points	Telescope
HE 0031-5525	2004-08-14 04:03	40	0.71	6464	1.6m
	2004-08-18 05:09	40	3.53	253	0.6m
	2004-08-19 04:00	30	4.80	418	0.6m
	2004-09-10 03:59	20	2.34	422	1.6m
	2004-09-11 02:31	30	0.88	105	1.6m
SDSS J024922.3-010006.7	2004-09-11 04:35	40	3.66	329	1.6m
	2004-11-13 01:56	20	2.44	439	1.6m
	2004-11-14 03:16	30	3.52	422	1.6m
	2004-11-15 00:30	15	2.56	615	1.6m
SDSS J125710.5+012422.9	2005-06-05 00:32	20	1.70	234	4.1m
	2005-06-07 00:03	30	2.00	193	4.1m
SDSS J153332.9-020600.0	2004-08-10 22:06	20	1.95	351	1.6m
	2004-08-11 21:56	20	1.81	325	1.6m
SDSS J161837.2-002302.7	2005-06-06 02:59	40	4.13	321	4.1m
	2005-06-07 05:54	40	2.00	160	4.1m
SDSS J164115.5+352140.6	2005-06-06 08:15	20	2.42	436	2.1m
	2005-06-09 07:14	30	0.40	48	2.1m
	2005-06-09 09:35	30	1.15	138	2.1m
SDSS J212808.4-000750.8	2004-08-11 00:20	30	2.15	258	1.6m
	2004-08-12 01:33	30	0.74	89	1.6m
SDSS J213530.3-074330.7	2004-08-11 02:43	30	1.71	205	1.6m
	2004-08-12 02:20	30	1.06	127	1.6m
SDSS J214723.7-001358.4	2005-07-05 08:14	40	2.10	115	4.1m
	2005-07-07 05:25	30	3.00	300	4.1m
SDSS J215354.1-073121.9	2005-06-06 07:33	30	2.00	200	4.1m
	2005-06-07 08:36	30	2.00	200	4.1m
SDSS J223135.7+134652.8	2004-08-13 01:34	40	3.88	349	1.6m
	2004-08-14 01:48	40	1.93	174	1.6m
	2004-08-18 01:26	60	1.58	82	0.6m
	2004-08-19 01:54	45	1.65	107	0.6m
	2004-09-10 00:30	30	3.14	377	1.6m
	2004-09-10 23:57	30	2.43	292	1.6m
SDSS J230726.6-084700.2	2004-08-11 04:32	40	1.28	115	1.6m

**Table 1.** Journal of observations with the 1.6-m and 0.6-m telescope at OPD, the 4.1-m SOAR telescope, and the 2.1-m telescope at McDonald observatory.  $\Delta T$  is the length of the observing run and  $t_{\text{exp}}$  is the integration time of each exposure.

Star	RA (2000)	DEC (2000)	$T_{\text{eff}}$ (K)	$\log g$	$g$ (mag)
HE 0031-5525	00:33:36	-55:08:37	11480 $\pm$ 30	7.65 $\pm$ 0.02	15.94
SDSS J024922.3-010006.7	02:49:22.3	-01:00:06.7	11060 $\pm$ 110	8.31 $\pm$ 0.10	19.08
SDSS J125710.5+012422.9	12:57:10.5	+01:24:22.9	11520 $\pm$ 160	8.36 $\pm$ 0.09	18.65
SDSS J153332.9-020600.0	15:33:32.9	-02:06:00.0	11350 $\pm$ 40	8.20 $\pm$ 0.02	16.62
SDSS J161837.2-002302.7	16:18:37.2	-00:23:02.7	10860 $\pm$ 160	8.16 $\pm$ 0.12	19.26
SDSS J164115.5+352140.6	16:41:15.5	+35:21:40.6	11230 $\pm$ 160	8.43 $\pm$ 0.10	19.04
SDSS J212808.4-000750.8	21:28:08.4	-00:07:50.8	11440 $\pm$ 100	8.29 $\pm$ 0.07	17.97
SDSS J213530.3-074330.7	21:35:30.3	-07:43:30.7	11190 $\pm$ 120	7.67 $\pm$ 0.09	18.59
SDSS J215354.1-073121.9	21:53:54.1	-07:31:21.9	11930 $\pm$ 130	8.07 $\pm$ 0.06	18.45
SDSS J223135.7+134652.8	22:31:35.7	+13:46:52.8	11080 $\pm$ 100	7.95 $\pm$ 0.07	18.63
SDSS J230726.6-084700.2	23:07:26.6	-08:47:00.2	11060 $\pm$ 110	8.19 $\pm$ 0.09	18.83

**Table 2.** List of properties of the new variables. The  $T_{\text{eff}}$  and  $\log g$  were derived from optical spectra, with the same model grid, despite different fitting procedures for HE and SDSS stars.

Star	$f$ ( $\mu\text{Hz}$ )	$P$ (s)	$A$ (mma)	$T_{\text{max}}$ (s)	$T_0$ (BCT)
HE 0031-5525	3035	329.5	$2.5 \pm 1.0$	$251 \pm 19$	2453231.6726750
	3611	276.9	$4.8 \pm 1.0$	$245 \pm 13$	
	3638	274.9	$1.5 \pm 1.0$	$40 \pm 41$	
SDSS J024922.3-010006.7	957	1045.2	$10.9 \pm 1.5$	$598 \pm 61$	2453259.6947745
	994	1005.6	$5.6 \pm 1.5$	$254 \pm 115$	
SDSS J125710.5+012411.9	1104	905.8	$46.7 \pm 5.7$	$31 \pm 18$	2453526.525547
SDSS J153332.9-020600.0	3837	260.6	$5.3 \pm 0.6$	$130 \pm 14$	2453228.4215328
	3879	257.8	$4.3 \pm 0.6$	$29 \pm 17$	
SDSS J161837.2-002302.7	1553	644.0	$5.4 \pm 1.1$	$238 \pm 22$	2453527.630436
SDSS J164115.5+352140.6	1276	809.3	$27.3 \pm 4.6$	$560 \pm 22$	2453527.750116
SDSS J212808.4-000750.8	3309	302.2	$17.1 \pm 1.3$	$52 \pm 4$	2453228.5198054
	3638	274.9	$11.0 \pm 1.3$	$154 \pm 6$	
	3461	289.0	$9.7 \pm 1.4$	$131 \pm 7$	
SDSS J213530.3-074330.7	1769	565.4	$49.8 \pm 2.3$	$12 \pm 5$	2453228.6189298
	1958	510.6	$16.8 \pm 2.2$	$18 \pm 14$	
	3094	323.2	$13.0 \pm 2.3$	$24 \pm 12$	
	3334	299.9	$22.9 \pm 2.3$	$209 \pm 6$	
	3549	281.8	$13.3 \pm 2.3$	$276 \pm 10$	
SDSS J215354.1-073121.9	4757	210.2	$5.6 \pm 0.9$	$15 \pm 5$	2453527.816768
SDSS J223135.7+134652.8	1413	707.5	$17.1 \pm 1.7$	$389 \pm 16$	2453230.5708122
	1595	627.0	$26.3 \pm 1.7$	$172 \pm 10$	
	1614	619.7	$18.9 \pm 1.7$	$104 \pm 13$	
	1822	548.7	$13.7 \pm 1.7$	$264 \pm 16$	
	2615	382.4	$14.6 \pm 1.7$	$369 \pm 10$	
SDSS J230726.6-084700.2	825	1212.2	$25.6 \pm 2.4$	$478 \pm 32$	2453228.6946388
	1621	617.0	$12.5 \pm 2.4$	$457 \pm 33$	

**Table 3.** Periodicities identified in our data sets for the new ZZ Ceti, where  $f$  is the frequency,  $P$  is the period, and  $A$  is the amplitude. The times of maxima ( $T_{\text{max}}$ ) are given in relation to the  $T_0$ , in Barycentric Coordinate Time in days.